









# Search for high frequency gravitational waves in electromagnetic cavities

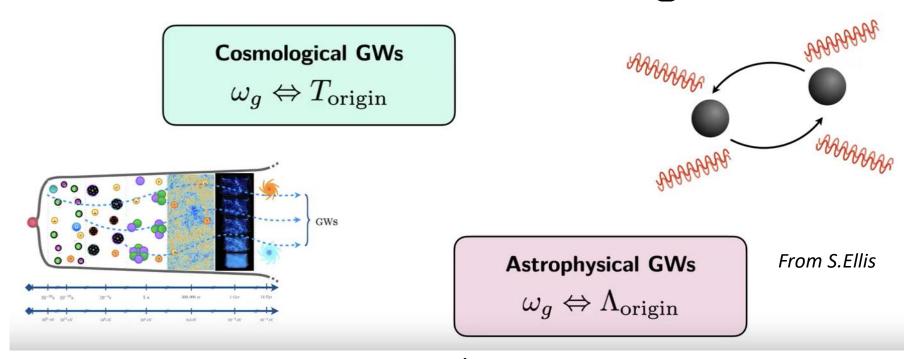
#### Jordan Gué

IFAE, Universitat Autonoma de Barcelona

In collaboration with T. Krokotsch (Universität Hamburg)

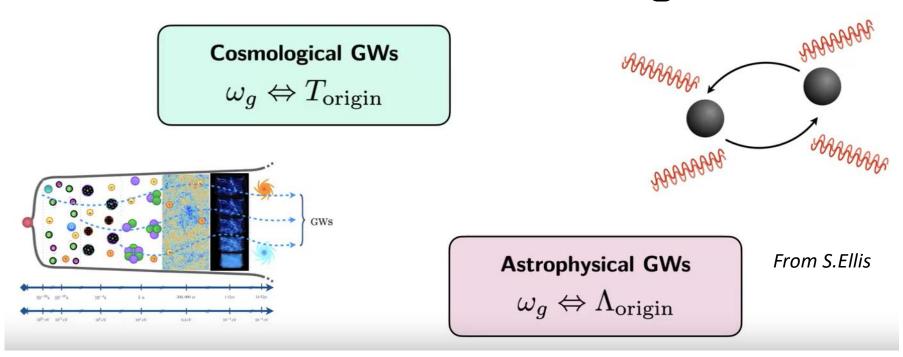
7<sup>th</sup> Barcelona Initiative for Gravitation

## HFGW astro/cosmo signals



→ Higher GW frequency ⇔ Higher energy scale/Lower length scale we can probe

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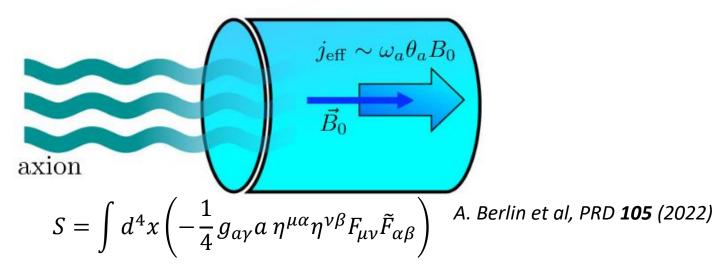
#### Some astrophysical sources:

- Mergers of PBH
- Mergers of ECO (boson stars,...)
- First order phase transition in neutron stars
- Superradiant boson clouds orbiting SMBH (monochromatic)

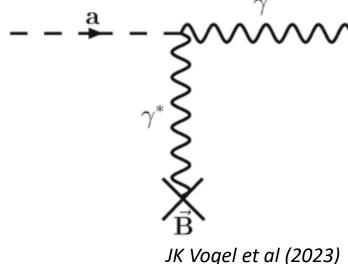
See N. Aggarwal et al, arXiv 2501.11723

→ Many of those sources are BSM

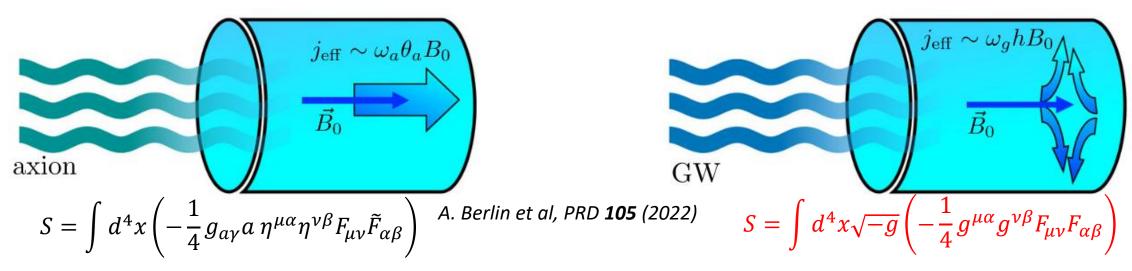
## Analogies with axion dark matter



- Simple way of looking for axions coupled to EM is through inverse Primakoff effect
- → Use of microwave cavities to search for GHz axions

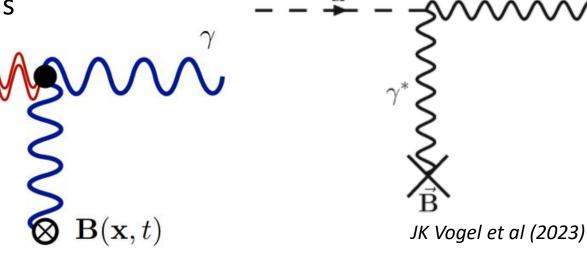


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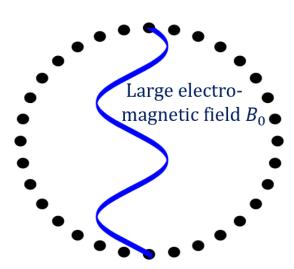
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- GW analog: inverse Gertsenshtein effect
- → Same apparatus is sensitive to HFGW



Credit: S. Ellis

# **Expected GW signals**

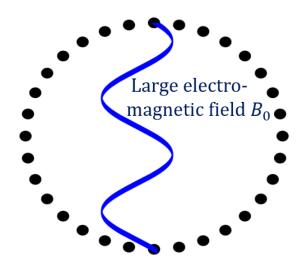


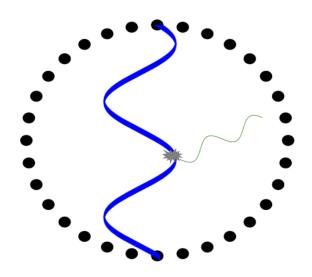
## **Expected GW signals**

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$$S = \int d^4x \sqrt{-g} \left( -\frac{1}{4} g^{\mu\alpha} g^{\nu\beta} F_{\mu\nu} F_{\alpha\beta} \right) \rightarrow \partial_{\nu} \delta F^{\mu\nu} \equiv j_{\rm eff}^{\mu} \propto \omega_g h B_0$$

→ GW couples to EM energy

A. Berlin et al, PRD **105** (2022) V. Domcke et al, PRL **129** (2022)





From T. Krokotsch

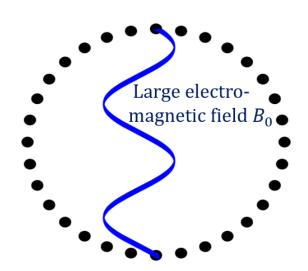
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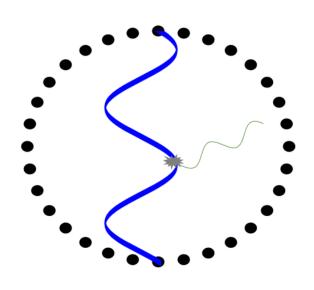
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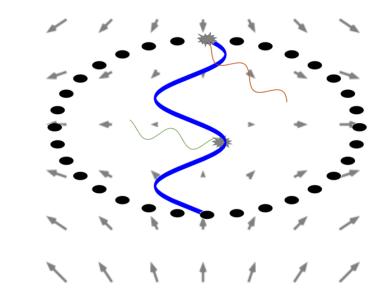
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- $\delta \ddot{x}_i \partial_j \sigma_{ij} = F_i^h$ M. Hudelist et al, CQG **40** (2023)
- → GW couples to mechanical energy



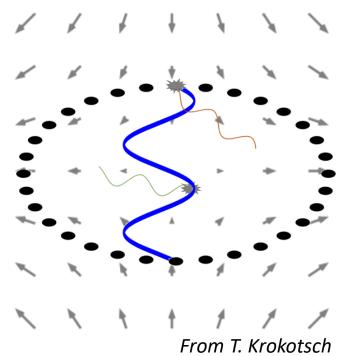




## Observables

- $\partial_{\nu} \delta F^{\mu\nu} = j_{\text{eff}}^{\mu}$
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In general, a dipole antenna measures  $\vec{E}$  at a point inside the cavity. How should we define  $\vec{E}$  ?



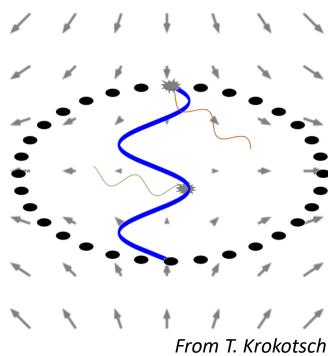
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- $\rightarrow F_{\alpha\beta}$  is covariant not invariant  $\rightarrow E_i^{obs} \neq F_{i0}$ . Instead,

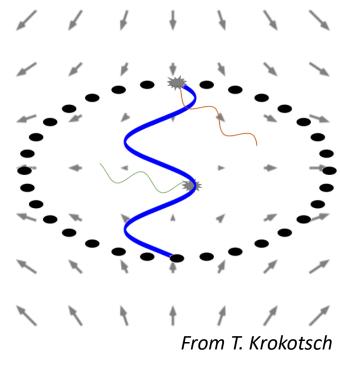


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$$E_{\underline{a}}^{obs} = F_{\mu\nu} u^{\nu} e_{\underline{a}}^{\mu} \qquad \qquad \text{Infinitesimal coord. system}$$

$$g_{\mu\nu} e_{\underline{a}}^{\mu} e_{\underline{b}}^{\nu} = \eta_{\underline{a}\underline{b}} \; ; e_{\underline{0}}^{\nu} = u^{\nu}$$

$$\Rightarrow \text{used to build a local Lorentz frame}$$
Observer's 4-velocity

Linearizing, 
$$\delta E_{\underline{a}}^{obs} = \delta F_{\mu\nu} e_{\underline{a}}^{\mu} u^{\nu} + F_{\mu\nu} \delta e_{\underline{a}}^{\mu} u^{\nu} + F_{\mu\nu} e_{\underline{a}}^{\mu} \delta u^{\nu} + \delta x^{\rho} (\partial_{\rho} F_{\mu\nu}) e_{\underline{a}}^{\mu} u^{\nu}$$

 $\rightarrow \delta F_{a0}$  is **not** the observed field in curved spacetime

## **Frames**

In general, to compute GW signals, choice between 2 frames: TT and PD

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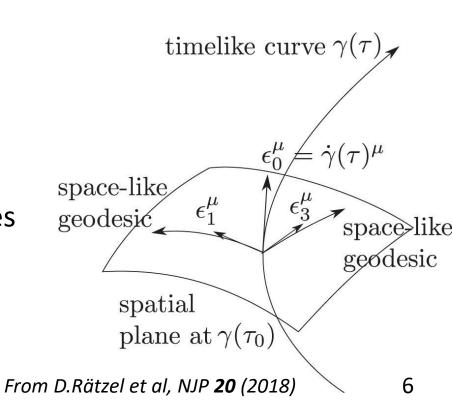
In general, to compute GW signals, choice between 2 frames: TT and PD

- Traceless-Transverse (TT) gauge : global coord. system set by freely falling test masses  $h_{0\mu}=\partial_i h^{ij}=h=0$
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- Proper Detector (PD) Frame: coordinate system built by extending observer's tetrads into geodesics
- → More intuitive : GW acts as a Newtonian force on rigid bodies
- → Metric perturbation more involved



#### Which frame should we use?

Freely falling limit:  $\omega_g \gg v_s/L$ 

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#### What about axion haloscopes?

In conductor,  $v_{\rm S}\sim 10^{-5}$ , i.e for a cavity with  $L\sim 0.1$  m, and  $\omega_{q}\sim {\rm GHz},~\omega_{q}L/v_{\rm S}\gg 1$ 

→ TT more convenient

$$\delta E_{\underline{a}}^{TT,FF} = \delta F_{a0} + F_{\mu 0} \delta e_{\underline{a}}^{\mu} + F_{av} \delta u^{v} + \delta x^{\rho} (\partial_{\rho} F_{a0}) = \delta F_{a0}$$

## Signal power

 $\rightarrow$  In TT, at high frequency,  $\partial_{\nu}\delta F^{\mu\nu}=j^{\mu}_{\rm eff}$  solved by expanding  $\delta F_{a0}$  in cavity eigenmodes.

On resonance, the signal power in a mode  $\vec{E}_n$  is given by  $P_{sig} = \frac{1}{2} Q \omega_g V \eta_g^2 h^2 \big| \vec{B} \big|^2$ 

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Adapted from A. Berlin et al, PRD 105 (2022)

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The signal power from axion DM is  $P_{sig}^a = \frac{1}{2}g_{a\gamma}^2Q\omega_aV\eta_a^2a^2\left|\vec{B}\right|^2$  with  $\eta_a = \frac{|\text{J }avE_n.B|}{\sqrt{V\int dV\left|\vec{E}_n\right|^2}}$ 

 $\rightarrow$  Up to  $\mathcal{O}(0.1)$  couplings, we have  $h=ag_{av}\sim 10^{-22}$ 

P. Sikivie, RMP 93 (2021)

Adapted from A. Berlin et al, PRD 105 (2022)

## Gauge invariance

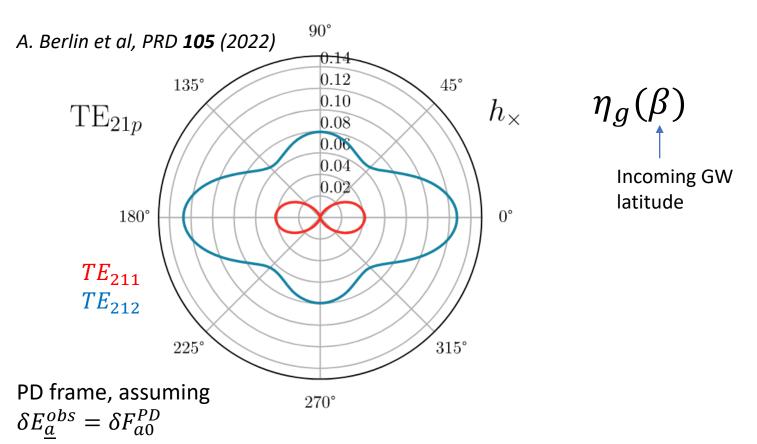
Considering  $\omega_g \sim \mathcal{O}(\text{GHz})$ , and background magnetostatic field, the observed electric field is

- In TT,  $\delta E_{\underline{a}}^{obs} = \delta F_{a0}^{TT}$  and  $\partial_{\nu} \delta F^{\mu\nu,TT} = j_{\mathrm{eff}}^{\mu,TT}$
- In PD,  $\delta E_{\underline{a}}^{obs} = \delta F_{a0}^{PD} + F_{ai} \delta u_i^{PD}$  and  $\partial_{\nu} \delta F^{\mu\nu,PD} = j_{\rm eff}^{\mu,PD}$

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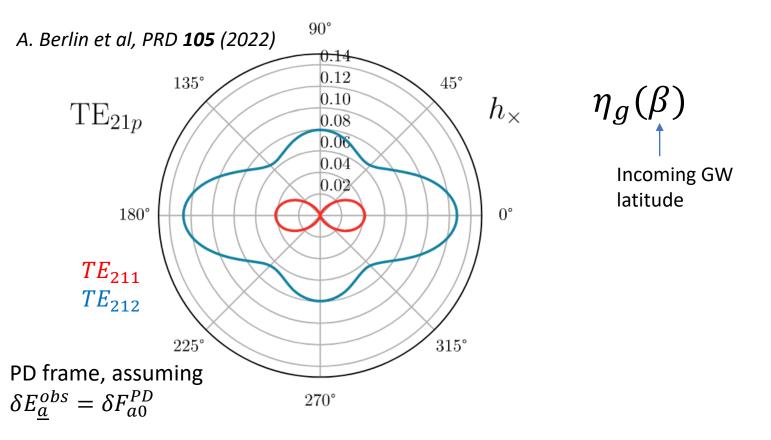
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T. Krokotsch, JG, in preparation



Mech. + EM (PD)

180°

Mech. (PD) EM (PD)

EM (TT)

90°

- EM signal  $\delta E_{\underline{a}} = \delta F_{a0} + F_{\mu 0} \delta e^{\mu}_{\underline{a}} + F_{a\nu} \delta u^{\nu} + \delta x^{\rho} (\partial_{\rho} F_{a0})$  and  $\partial_{\nu} \delta F^{\mu \nu} = j^{\mu}_{\rm eff}$ 
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- Heterodyne/Homodyne setups

### Conclusion

• Microwave cavities are powerful probes of monochromatic HFGW ( $h \sim 10^{-22}$ )

GW couples to all types of energy,
 care must be taken to model all effects

- With current quantum technology, this is not enough to probe cosmological GW
- → Cross correlate multiple cavities : GravNet ERC
- → Use of Earth modulation for persistent signals
- → Quantum enhancement techniques (e.g. squeezing)

